

## SPECKLE SUPPRESSION WITH THE PROJECT 1640 INTEGRAL FIELD SPECTROGRAPH

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### ABSTRACT

Project 1640 is a high-contrast imaging instrument recently commissioned at the Palomar observatory. A combination of a coronagraph with an integral-field spectrograph (IFS), Project 1640 is designed to detect and characterize extrasolar planets, brown dwarfs, and circumstellar material orbiting nearby stars. In this paper, we present our data processing techniques for improving upon instrument raw sensitivity via the removal of quasi-static speckles. Our approach utilizes the chromatic image diversity provided by the IFS in combination with the locally optimized combination of images algorithm to suppress the intensity of residual contaminating light in close angular proximity to target stars. We describe the Project 1640 speckle suppression pipeline and demonstrate its ability to detect companions with brightness comparable to and below that of initial speckle intensities using on-sky commissioning data. Our preliminary results indicate that suppression factors of at least one order of magnitude are consistently possible, reaching  $5\sigma$  contrast levels of  $2.1 \times 10^{-5}$  at  $1''$  in the  $H$  band in 20 minutes of on-source integration time when non-common-path errors are reasonably well calibrated. These results suggest that near-infrared contrast levels of order  $\approx 10^{-7}$  at subarcsecond separations will soon be possible for Project 1640 and similarly designed instruments that receive a diffraction-limited beam corrected by adaptive optics systems employing deformable mirrors with high actuator density.

*Key words:* instrumentation: adaptive optics – techniques: high angular resolution – techniques: spectroscopic

*Online-only material:* color figures

### 1. INTRODUCTION

Wave-front errors created by imperfections in the optical components (mirrors, lenses, beam splitters, etc.) of a high-contrast imaging instrument manifest as a complex pattern of quasi-static intensity variations, or so-called speckles, in the focal plane at the science camera. Nominally bright, compared to the signal of a substellar companion or debris disk, speckle intensities must first be minimized by hardware, usually with an adaptive optics (AO) system, and then further suppressed through data processing. Since individual speckles have the same intrinsic width as the instrument point-spread function (PSF), this latter step is crucial for distinguishing between stellar artifacts and true companions.

Speckle suppression is accomplished through differential imaging by exploiting differences between the properties of residual starlight that reaches the detector and any incoherent radiation arriving from an off-axis source. To preserve companion flux, some form of “diversity” must be introduced into the system, whereby a sequence of images is recorded such that individual frames are trivially different from one another. Once measured or calibrated, the image diversity is reversed using software to remove the speckles and reveal previously hidden companions.

Image diversity can be achieved in a number of ways. Examples include modulation of companion position with respect to speckles through instrument rotation or field rotation (Krist

2007; Marois et al. 2008, 2010b; Leconte et al. 2010), modulation of input polarization state (Potter 2003; Oppenheimer et al. 2008; Hinkley et al. 2009), and modulation of the actual target being observed (Serabyn et al. 2010; Crepp et al. 2010; Mawet et al. 2009). The level to which residual starlight is removed depends on the modulation timescale compared to the speckle decorrelation timescale. Quasi-static speckles<sup>9</sup> generally decorrelate in tens or hundreds of seconds as the result of changes in ambient conditions and optical system alignment (e.g., Hinkley et al. 2007). Ideally, science images and PSF reference images are recorded simultaneously to minimize the influence of speckle pattern evolution.

An alternative approach to discriminate between companions and starlight is to generate diversity through the wavelength dependence of speckles. In this case, modulation of companion position with respect to speckles is achieved naturally through diffraction by acquiring images simultaneously in multiple different spectral channels. Chromatic differential imaging not only has the benefit of “freezing” the speckle pattern in time, but it also yields the spectrum of any companions, thus enabling study of their chemical composition and bulk physical properties. First proposed by Sparks & Ford (2002), concurrent search and characterization may be accomplished in practice using an integral-field spectrograph (IFS).

<sup>9</sup> Other families of speckles related to the atmosphere and adaptive optics system can decorrelate on a much faster timescale (Macintosh et al. 2005).

Compared to other speckle suppression techniques (Absil & Mawet 2010; Oppenheimer & Hinkley 2009), chromatic differential imaging has a high duty-cycle efficiency and can help maximize instrument scientific return by: accommodating a relatively wide bandpass<sup>10</sup>; allowing for the inter-changeable usage of science images and PSF references; maintaining an inner-working-angle that is independent of target coordinates relative to the observatory; avoiding smearing of the companion or dust disk PSF during an exposure; and obviating the need to observe a nearby PSF calibration star with similar brightness and colors. Chromatic differential imaging may also be implemented in combination with other techniques.

AO-assisted integral field spectroscopy has been used previously to study the companions orbiting GQ Lupi (McElwain et al. 2007), AB Doradus (Thatte et al. 2007), AB Pictoris (Bonnetfoy et al. 2010), and recently the outer planets of HR 8799 (Bowler et al. 2010). In this paper, we present the first on-sky experiments that combine an IFS with a coronagraph (Hinkley et al. 2011) and speckle suppression using chromatic differential imaging (Sparks & Ford 2002). To this, we also implement the locally optimized combination of images (LOCI) algorithm which improves the signal-to-noise ratio of detections (Lafrenière et al. 2007a).

Project 1640 (hereafter, P1640) is a ground-based high-contrast imaging instrument that was recently installed and tested on the Hale 200 inch telescope at Palomar. The hardware incorporates a near-infrared coronagraph and IFS and will soon receive a corrected beam from the PALM-3000 “extreme” AO system (Hinkley et al. 2011; Bouchez et al. 2010; Soumerai 2005). P1640 has made several discoveries to date, including Alcor B and Zeta Virginis B (Zimmerman et al. 2010; Hinkley et al. 2010). Though faint compared to their host stars, these companions have masses of  $\approx 0.25 M_{\odot}$  and  $\approx 0.17 M_{\odot}$ , respectively, and were noticed in raw data.

In the following, we describe our technique for detecting companions having brightness and angular separation that places them beneath the noise floor set by speckles prior to data processing. This paper is accompanied by a companion paper by L. Pueyo et al. (2011, in preparation) that discusses our method for subsequently extracting their spectrum. Combined, our results demonstrate the two principal utilities of using an IFS for high-contrast observations. These techniques are relevant to forthcoming instruments with similar designs, including the Gemini Planet Imager (Macintosh et al. 2006) and VLT SPHERE (Beuzit et al. 2006), which will also use an IFS for chromatic differential imaging and companion characterization.

## 2. DATA PIPELINE DESCRIPTION

### 2.1. IFS Raw Data

The P1640 IFS records 40,000 different spectra across the instrument field of view (FOV) with every exposure. Each spectrum corresponds to an individual element of a  $200 \times 200$  lenslet array and is mapped to a different location in the image plane. Each lens in the array has a pitch of  $\approx 19.2$  mas as projected onto the sky. The P1640 spectral resolution is 23 channels across the *J* and *H* bands, spanning wavelengths  $1.10 \mu\text{m} \leq \lambda \leq 1.75 \mu\text{m}$  (Hinkley et al. 2011). Raw spectra are converted into a data cube consisting of two spatial coordinates

<sup>10</sup> A special case of chromatic differential imaging, called simultaneous differential imaging (SDI), operates over a narrow bandpass and has shown promise for high-contrast applications involving methanated sources as demonstrated by the NICI campaign (Billler et al. 2010; Liu et al. 2010).

and one color coordinate using the procedure described in detail in Zimmerman et al. (2011). Each extracted cube is flat-fielded, background-subtracted, and wavelength-calibrated based on observations of spectral standards. The P1640 speckle suppression pipeline (PSSP) starts with extracted and spectrally calibrated data cubes.

### 2.2. The PSSP Code

PSSP is a custom program written in Matlab. This language was selected because it is flexible and contains a number of useful, fast, and robust built-in functions. Moreover, Matlab allows for a trivial transition from running scripts in serial on one CPU core to running the same scripts in parallel on multiple cores, facilitating handling of the large number of images and calculations per image. To enhance the reduction speed, we used the Bluedot super-computing cluster at the NASA Exoplanet Science Institute (NExSci) which incorporates 16 nodes with 8 cores (2.5 GHz) per node. This feature permits pseudo-real-time reduction while at the observatory, which enables follow-up observations of promising targets on subsequent nights within the same run to improve the fidelity of detections, spectra, and astrometry, and possibly take advantage of improved seeing conditions.

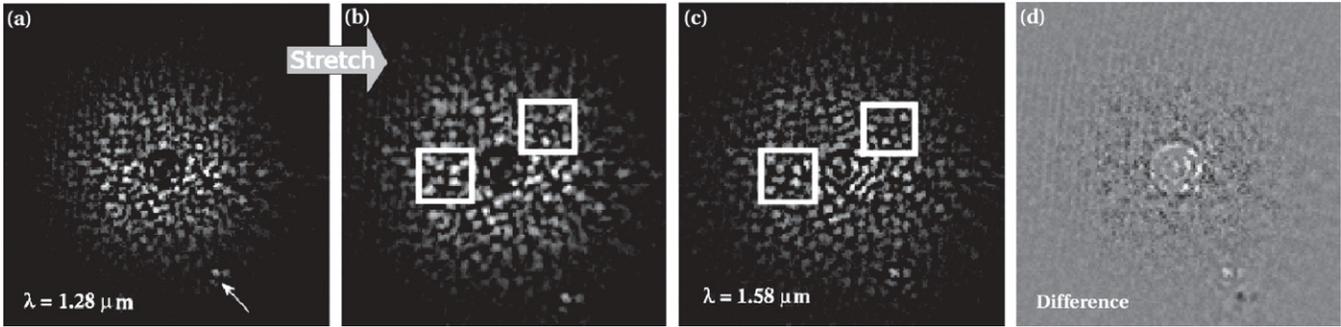
### 2.3. Technique

At its core, PSSP involves precision stretching and shrinking of individual images to utilize the spatial dependence of the speckle pattern on wavelength. This technique takes advantage of the fact that the position of a true companion is constant in each channel while speckles move radially outward from the star as a function of wavelength. For instance, the dependence of speckle position in the focal plane scales linearly with wavelength for phase aberrations created in the pupil plane. This form of diversity lends itself to differential imaging (Section 1). Stretching and shrinking images separated by several wavelength channels simultaneously aligns speckles and misaligns companions (Figure 1). Upon subtraction, the highly correlated speckles are removed and companions survive, dramatically improving the effective detection sensitivity (Section 3).

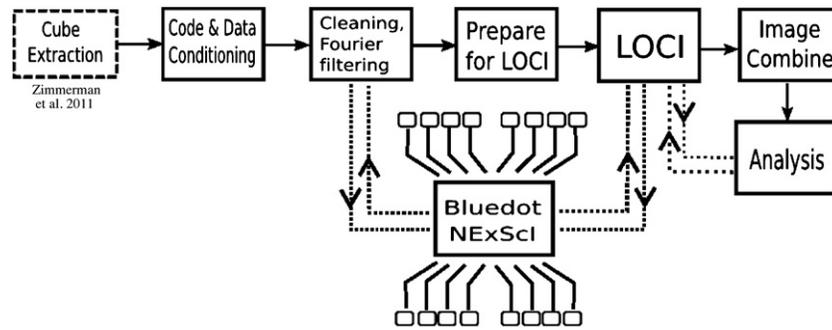
The multitude of available reference images provided by the IFS naturally lends itself to application of the LOCI algorithm (Lafrenière et al. 2007a). Once images are optimally stretched for a given wavelength channel, we apply LOCI to construct an optimal reference, including both color and time, by weighting individual images with coefficients based on a least-squares matching of local speckle intensities via matrix inversion. LOCI has been used with other speckle suppression techniques and provides higher signal-to-noise ratio detections compared to PSF subtraction involving a single reference image (Marois et al. 2008; Mawet et al. 2009; Lafrenière et al. 2009; Crepp et al. 2010). Details of the algorithm are discussed below.

### 2.4. Program Outline

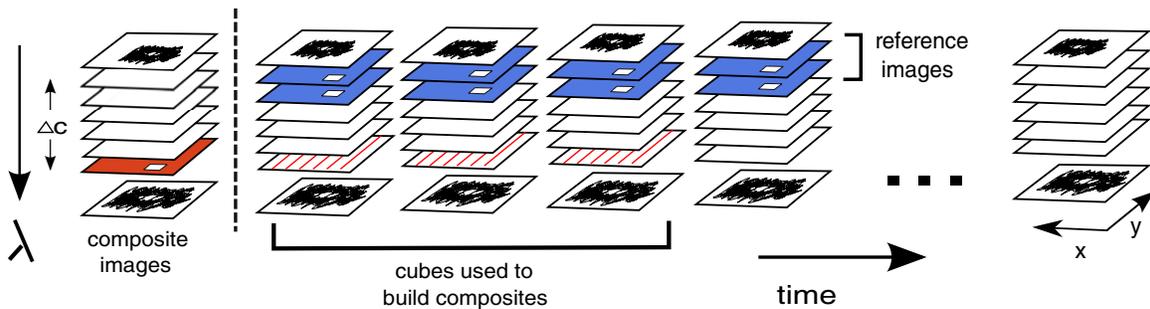
An outline of the PSSP procedure is shown in Figure 2. We begin by retrieving extracted data cubes output by the Zimmerman et al. (2011) raw IFS conversion program. Data headers are then culled for pertinent information and the code is preconditioned by the user. This latter step involves: specifying which images to use based on data quality; searching for



**Figure 1.** Occulted images of the star Alcor in individual IFS channels illustrating differential imaging with chromatic speckle diversity. Each image has a width of  $3'' \times 3''$ . (a) Single channel in the *J* band after applying a high-pass Fourier filter but prior to image stretching. The position of Alcor B is marked with an arrow. (b) Same image radially stretched to optimally match the speckle pattern for a longer spectral channel. (c) Single channel in the *H* band. The speckle structure is highly correlated between the middle panels (see the interior of white boxes), enabling precision PSF alignment and subtraction. (d) Differenced image displayed on a linear intensity scale. The signal of a real companion is incoherent with starlight and survives post-processing because its position is independent of wavelength. Non-differenced images are displayed on a square-root intensity scale for comparison with the speckle patterns.



**Figure 2.** Block diagram of the P1640 speckle suppression pipeline (PSSP). Images are first extracted from raw IFS data via the procedure described in Zimmerman et al. (2011) to form cubes. The cubes are then processed using wavelength diversity and the LOCI algorithm to detect faint companions that are originally buried in speckle noise. Multiple CPU nodes from the NASA Exoplanet Science Institute (NExSci) Bluedot super-computer run custom Matlab scripts in parallel to enhance speed at several stages.



**Figure 3.** Schematic of a P1640 hyper cube. Composite images shown to the left of the vertical dashed line are built by co-adding frames in each wavelength channel across the time domain to increase companion signal-to-noise ratio prior to performing PSF subtraction. In this example, images with red stripes are used to create the image in solid red. An optimal PSF reference image for each composite is constructed from the many available reference frames in the color and time domains (solid blue images). A small sub-frame is shown to indicate that PSF subtraction occurs locally. Composite images and usable reference frames are separated by  $\Delta c$  wavelength channels to minimize subtraction of companion signal.

(A color version of this figure is available in the online journal.)

astrometric fiducial spots<sup>11</sup>; generating FOV, high-pass Fourier filter, and other alignment masks; setting basic parameters related to image cleaning and the LOCI algorithm; inserting fake companions to assess the fidelity of the reduction; and initializing the number of Matlab “workers” for parallel processing. Both occulted and unocculted images are loaded in order to perform relative photometry and calculate sensitivity in terms of contrast.

Depending on target near-infrared brightness, several dozen cubes are generated for each star. This large data set, which we refer to as a hyper-cube when the time dimension is included (Figure 3), consists of hundreds of images. In each channel of each cube, the spatial pixels are cleaned using standard techniques to correct for any remaining intensity outliers and negative values. Depending on whether we are searching for point sources or diffuse circumstellar material, like a dust disk, the images may be high-pass Fourier filtered to remove the low-frequency component of the stellar halo and enhance the signal of both companions and speckles—the latter of which is useful for image alignment. The low-frequency component

<sup>11</sup> A grid of wires placed in the pupil plane diffracts starlight into four bright spots centered on the star and may be used for astrometric purposes (Sivaramakrishnan & Oppenheimer 2006; Zimmerman et al. 2010).

of the image is removed using a quadratic attenuation profile, which we find has better noise properties than a linear or step-function filter. This stage of the procedure is time consuming and so is sent to the NExScI Bluedot cluster for processing.

Once cleaned, the images in each wavelength channel are stacked across the hyper-cube time dimension to form composite images (Figure 3). This step increases the companion signal-to-noise ratio in preparation for the LOCI algorithm. The speckles themselves are used to register images. We find that this approach is efficient and as reliable as using the astrometric grid spots for alignment, provided a custom gray-scale mask is used to prevent the coronagraphic spot or companion itself from biasing the result. A Fourier-based program registers the images with 0.01 pixel precision (Guizar-Sicairos et al. 2008). We find that allowing for small distortions in the position of the speckles, by treating local image regions as a flexible membrane, results in only a marginal improvement in PSF matching. The number of temporal cubes used to form composite images depends on data quality and quantity. In general, cubes with sufficiently high speckle signal-to-noise ratio may be divided into a larger number of groups that are treated separately by the PSSP. The results may then be combined following image subtraction.

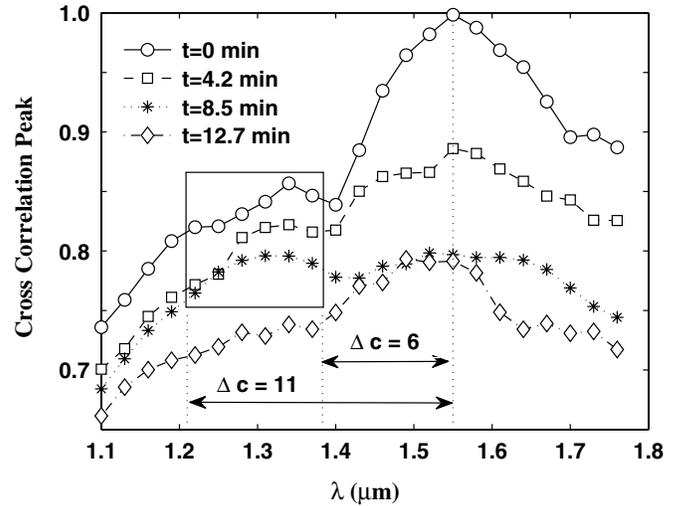
The LOCI algorithm constructs an optimal PSF reference for each wavelength channel of the composite cube using neighboring frames in color and time. To generate sufficient wavelength diversity, and therefore minimize subtraction of companion light, the reference frames associated with a given composite image must be separated by several spectral channels,  $\Delta c = c_0 - c_i$ , where  $c_0$  is the channel of the composite image with wavelength  $\lambda = \lambda_0$  and  $c_i$  is the channel common to a given set of reference frames with wavelength  $\lambda = \lambda_i$ . The minimum difference between channels is a function of composite image wavelength,  $\lambda_0$ , spectral resolution,  $\delta\lambda$ , and field angle,  $\theta = \theta_0$ ,

$$|\Delta c_{\min}| \approx \frac{\alpha \lambda_0}{\theta_0 \delta\lambda}, \quad (1)$$

where  $\alpha \approx 1$  is the fraction of a diffraction width that PSF reference images are radially stretched and  $\theta_0$  is expressed in the same units (unitless). The optimal  $\Delta c$  range is determined by the competing effects of companion throughput and speckle correlation between channels.

Figure 4 displays the degree of correlation between images as a function of wavelength and time following cleaning and optimal stretching and shrinking. All images are compared to a channel in the *H* band corresponding to a wavelength of  $\lambda_0 = 1.55 \mu\text{m}$  in the first cube. A box is drawn around the reference images that may be sent to the LOCI algorithm for this case. The difference between two sets of wavelength channels is labeled. We generally include channels in the water bands, located around  $1.35 \mu\text{m} \lesssim \lambda_i \lesssim 1.50 \mu\text{m}$ , despite their comparatively lower signal due to atmospheric absorption. It is likewise possible to use longer wavelengths to achieve chromatic diversity. This figure demonstrates the utility of an IFS for speckle suppression. Speckles in stretched images are highly correlated in color—the degree to which is limited only by the intrinsic wavelength dependence of the instrument—but decorrelate after several minutes in time.

Reference images that are closely spaced in time with  $\Delta c \geq \Delta c_{\min}$  are optimally stretched using an iterative least-squares technique for speckle spatial matching that incorporates precision registration via the Fourier-based method mentioned above. Images are resampled using bicubic interpolation, where the intensity of an output pixel is calculated from a weighted

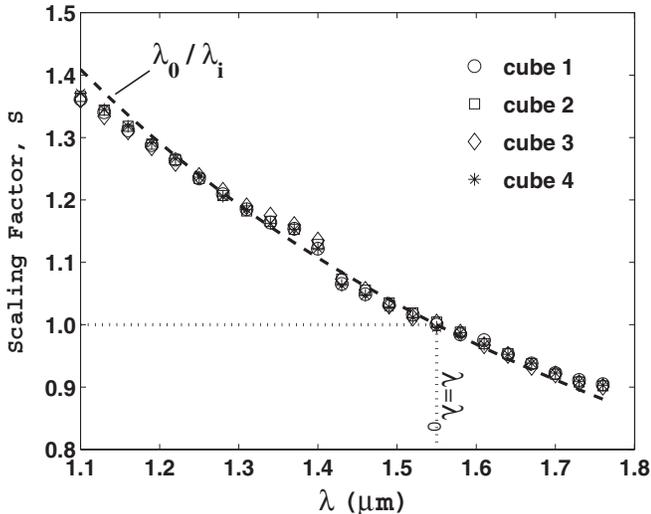


**Figure 4.** Cross-correlation peaks in color and time for the star Alcor after optimal image stretching. The star is occulted by the coronagraph, the cubes have been high-pass Fourier filtered, and Alcor B has been masked. All images are compared to the  $\lambda_0 = 1.55 \mu\text{m}$  channel in the first cube. The legend shows the amount of time elapsed in minutes between exposures. Speckles are highly correlated as a function of wavelength (see also Figure 1), thus justifying the use of processed IFS data for PSF reference subtraction. When constructing an optimal reference image via the LOCI algorithm, several channels must be skipped (given by  $\Delta c$ ) to prevent subtraction of any companions. We nominally use cubes acquired within a certain interval of time. These images are identified as enclosed by a box. Different channels have different reference images. In this example, there are 18 highly correlated images available that provide sufficient wavelength diversity for the  $1.55 \mu\text{m}$  channel. More images may be used in general since the LOCI algorithm will assign their associated coefficients an appropriate weight.

average of the neighboring  $4 \times 4$  pixels. A first guess for the optimal stretching parameter,  $S$ , is based on the wavelength of channels of interest using the scaling relationship,  $S = \lambda_0/\lambda_i$ , expected from speckle movements for phase aberrations located solely in the pupil plane. Here,  $\lambda_0$  again represents the wavelength of a non-reference (composite) image and  $\lambda_i$  represents the wavelength of a reference image. From this starting point, we then adjust the scaling factor to converge to an optimal value accurate to  $\Delta S \approx 10^{-3}S$ .

Figure 5 shows the optimal scaling factors,  $S$ , plotted as a function of wavelength for the same reference channel and data set shown in Figure 4. The theoretical relationship,  $S = \lambda_0/\lambda_i$ , is overplotted for comparison. With each cube, we find that optimal scaling factors are systematically smaller than the expected  $\lambda_0/\lambda_i$  relation when  $\lambda_i < \lambda_0$ , but larger than  $\lambda_0/\lambda_i$  when  $\lambda_i > \lambda_0$ . This effect is understood intuitively by considering the radial displacement of speckles as a function of wavelength. The scaling relation,  $S = \lambda_0/\lambda_i$ , is derived for phase errors located in a pupil plane. In practice, wave-front phase errors are also generated by out-of-pupil-plane optics and reflectivity and transmission non-uniformities which can transform into phase aberrations (Shaklan & Green 2006). Consequently, the radial position of a given speckle,  $\theta(\lambda)$ , increments more slowly than the linear dependence expected from first-order diffraction theory and instead follows  $\theta(\lambda) \lesssim \lambda\theta_0$  as a function of wavelength.

Exceptions occur in the water bands ( $\lambda \approx 1.4 \mu\text{m}$ ) as a result of the cube extraction algorithm operating with low signal. In this regime, the spectrum in each spatial pixel becomes biased toward the red end of the *J* band and also blue end of the *H* band (see Zimmerman et al. 2011 for details). As a result, the cube extraction algorithm creates a spatial pattern that systematically



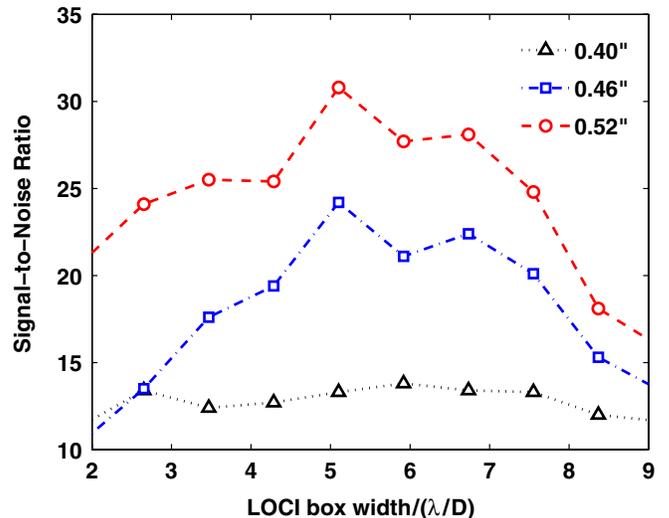
**Figure 5.** Image scaling (stretching/shrinking) factors as a function of wavelength for all IFS channels calculated for  $\lambda_0 = 1.55 \mu\text{m}$  using the same cubes of the star Alcor as Figure 4. If all aberrations where phase errors that occurred in the pupil plane, the optimal scaling factors would follow  $S = \lambda_0/\lambda_i$ , shown by the dashed line. Since aberrations are also generated in amplitude and by out-of-pupil-plane optics,  $S < \lambda_0/\lambda_i$  when  $\lambda_i < \lambda_0$ , and  $S > \lambda_0/\lambda_i$  when  $\lambda_i > \lambda_0$  (see the discussion). Image channels located between the  $J$ - and  $H$ -filters, near  $\lambda_i \approx 1.35\text{--}1.50 \mu\text{m}$ , have a lower signal as they are affected by atmospheric water bands, leading to a systematic offset and larger dispersion in the values by which images are scaled as a consequence of the cube extraction algorithm. We nominally use channels located in the water bands to build PSF reference images for bright stars. Optimal scaling factors are therefore calculated on an individual basis.

mimics an adjacent spectral channel, thus creating the bump seen in Figure 5. This systematic error does not effect PSF subtraction, but it does artificially modify the spectrum of companions in the water bands, requiring calibration. The standard deviation value of  $S$  averaged across the spectral bands is 0.003, comparable to image alignment precision. As a result of these effects, the PSSP recalculates the optimal scaling factor for each image in each cube for each star to maximize performance.

Our implementation of the LOCI algorithm operates on each spatial pixel in turn. This step is likewise time consuming and so is also sent to the Bluedot cluster. Local image coefficients are generally assigned based on the flux inside of a box of width  $5 \times 5$  PSF widths. Figure 6 shows optimization of the signal-to-noise ratio as a function of the LOCI box width (referred to as the O-zone in Lafrenière et al. 2007a) for three fake companions inserted into the Alcor data set. Their angular separations from the star are  $0''.40$ ,  $0''.46$ ,  $0''.52$ , respectively, as shown in Figure 7 and discussed in more detail below. Once complete, the program median-combines the reduced data in the color dimension. The  $J$  and  $H$  bands are separated at  $\lambda = 1.40 \mu\text{m}$  to avoid unnecessary averaging in case a companion is particularly red or blue.

### 3. ON-SKY DEMONSTRATION

P1640 has been used on several observing runs at Palomar in recent semesters. In this section, we present our preliminary on-sky results, demonstrating the ability to recover real and fake companions that are faint compared to the intensity of quasi-static speckles. We characterize the signal-to-noise ratio of detections and show our sensitivity in terms of contrast.



**Figure 6.** Signal-to-noise ratio vs. LOCI (outer) box width for three fake companions injected into the Alcor data set shown in Figure 7. The optimal LOCI box has a width of  $\approx 5 \lambda/D$ , reaching a compromise between speckle suppression and companion throughput. The companions have equal brightness and separations of  $0''.40$ ,  $0''.46$ , and  $0''.52$ . More wavelength diversity ( $\Delta c > 6$ ) is required to detect companions with separation  $< 0''.4$ .

(A color version of this figure is available in the online journal.)

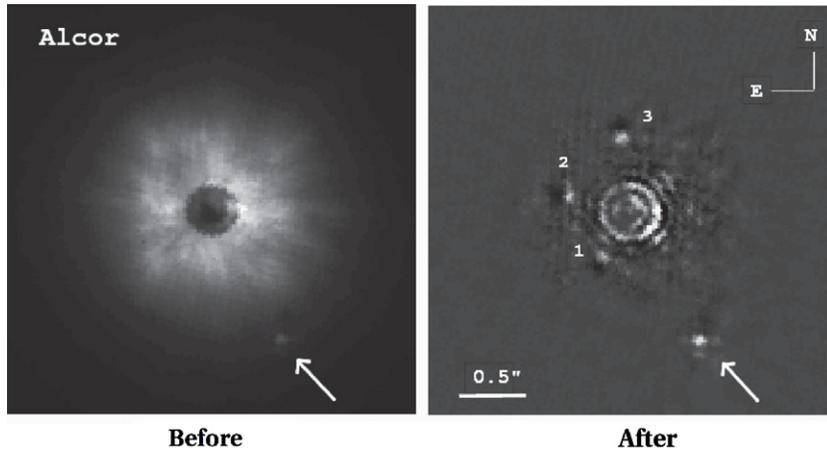
#### 3.1. Alcor

P1640 discovered an M-dwarf orbiting the nearby, bright A-star Alcor in 2009 March (Zimmerman et al. 2010). With an angular separation of  $\approx 1''$  and brightness ratio of  $\sim 6$  mag in  $JH$ , the  $\approx 0.25 M_\odot$  companion was identified in stacked cubes prior to speckle suppression.<sup>12</sup> In this section, we use the Alcor data set to demonstrate the ability to detect fake companions that are fainter and located closer to the primary star where quasi-static speckles are bright as a result of non-common-path wave-front errors.

To test PSSP and also calculations of contrast (Section 3.3), we inserted artificial companions of various brightness and angular separation into the Alcor on-sky data. The companions were given PSF widths and relative intensities that match Alcor B in each channel. To minimize PSF smearing in color and time, each companion was injected into each image channel of each cube by using Alcor B as an astrometric reference point. An optimal LOCI (outer) box size of  $5 \times 5$  PSF widths was used to maximize the signal-to-noise ratio (Figure 6). Figure 7 shows  $H$ -band images of the star before and after the application of PSSP using three artificial companions that are each 1 magnitude fainter than Alcor B with separations of  $0''.40$ ,  $0''.46$ , and  $0''.52$ , respectively.

The three fake companions are completely buried in speckle noise in pre-processed cubes and detectable only after using the wavelength diversity provided by the IFS. After application of PSSP, each display the tell-tale dipole pattern—a closely spaced positive and negative image—as a result of reference image stretching and subtraction. The inclusion of both image stretching and shrinking would yield a tripole pattern (Sparks & Ford 2002). It is possible to further build companion flux by subtracting composite images from reference images and then rescaling and stacking the results. This too creates a tripole

<sup>12</sup> Extracted IFS data cubes may be converted into a movie that sequentially displays images in wavelength space, allowing for the visual identification of candidate companions whose brightness is comparable to the intensity of speckles.



**Figure 7.** Stacked  $H$ -band images of the star Alcor before (left) and after (right) running the PSSP. Three fake companions have been injected into the on-sky data with separations of  $0''.40$ ,  $0''.46$ , and  $0''.52$ . The companions are 1 mag fainter than Alcor B in each channel and detectable only after using the wavelength diversity provided by the IFS. The dipole pattern is evident for the inner companions as a consequence of image stretching. Alcor B does not exhibit the dipole pattern because it is artificially masked in PSF reference images to facilitate image registration. The bright ring surrounding the coronagraph occulting spot results from insufficient wavelength diversity as a consequence of the slower outward movement of speckles with spectral channel in the regions very close to the star.

pattern, but it also smears the PSF and can degrade astrometric precision. Alcor B does not exhibit the dipole pattern because it is masked in each PSF reference image to facilitate precision image registration. The fake companions have a final signal-to-noise ratio of 13, 24, and 31 for respective separations of  $\theta = 0''.40$ ,  $0''.46$ , and  $0''.52$  as shown Figure 6. Alcor B is unambiguously recovered with a signal-to-noise ratio gain from 11.9 to 113.4.

The fake companions have a PSSP throughput of 42% (inner), 35% (middle), and 51% (outer), where throughput is defined as the ratio of the total companion flux before and after data processing averaged over each wavelength channel in the  $H$  band. Throughput is a function of the local speckle noise and angular separation and must be calibrated to accurately recover the spectro-photometric signal of companions that are fainter than stellar artifacts prior to applying PSSP. We find that broadband aperture photometry measurements typically result in 10% errors following calibration, with larger uncertainties occurring closer to the star. More precise measurements may be obtained by selecting PSF reference images that minimize fluctuations in the speckle subtraction residuals, but even 5% photometry is challenging given the many inherent biases present in LOCI data sets that must be accounted for (e.g., Marois et al. 2010a). The topic of precise spectral extraction is discussed in detail in L. Pueyo et al. (2011, in preparation).

No real companions closer to the Alcor primary were detected. We find that fake companions injected with a yet fainter signal, by factors of  $\approx 2.6$ , 4.8, and 6.2 at separations of  $\theta = 0''.40$ ,  $0''.46$ , and  $0''.52$ , are recoverable, respectively. This result is consistent with Figure 6 and was also used to check PSSP contrast calculations (see Section 3.3). Alcor is a member of the Ursa Majoris moving group which has a purported age of 400–600 Myr (Castellani et al. 2002; King et al. 2003). Using the Zimmerman et al. (2010)  $H$ -band photometry of the primary and secondary along with our contrast measurements, we are able to rule out additional companions with mass  $m \geq 77 M_J$  and separation  $\theta \geq 0''.52$  (Baraffe et al. 2003).

### 3.2. *FU Orionis*

FU Orionis (Ori) is a prototypical young stellar object after which a class of enigmatic low-mass stars exhibiting signs

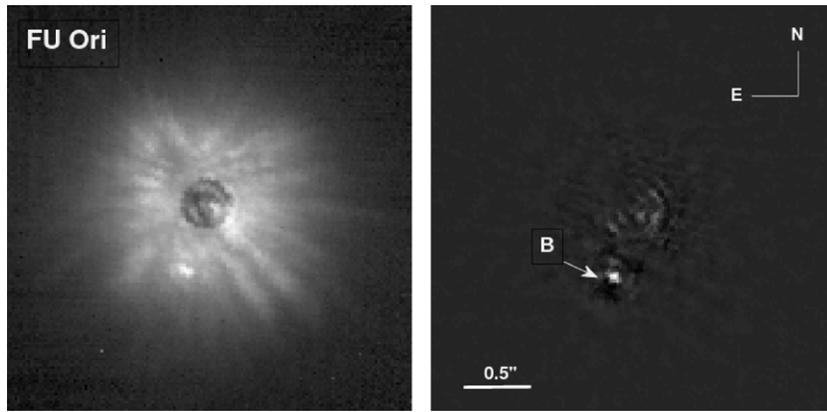
of disk accretion is named. Such objects experience outbursts resulting in 4–5 mag of visual brightening on timescales of months, with decay over decades sometimes punctuated by smaller scale episodes (Hartmann & Kenyon 1996). FU Ori itself has a known stellar companion separated by  $0''.5$  (Wang et al. 2004). It is thought that many FU Ori objects are members of binary or hierarchical triple systems (Reipurth & Aspin 2004).

We observed FU Ori on 2009 March 17 to better characterize the companion, which we measure to be 5.2 mag fainter than the primary in the  $J$  band and 4.4 mag fainter in the  $H$  band. Given the level of contamination, FU Ori B does not currently have a published  $J$ -band spectrum as it is mixed with speckle noise. Indeed, it was originally discovered using PSF subtraction of a nearby calibration star (Wang et al. 2004).

Figure 8 shows combined  $JH$  images of FU Ori before and after running the PSSP algorithm. The companion is just perceptible in pre-processed data. Once applied, PSSP unambiguously recovers the companion as indicated by the bright positive and negative signature. Analysis of individual channels shows that we are able to improve from a  $1.5\sigma$  to  $16.4\sigma$  detection at  $\lambda = 1.22 \mu\text{m}$  and  $2.3\sigma$  to  $22.9\sigma$  detection at  $\lambda = 1.58 \mu\text{m}$ , where  $\sigma$  is the standard deviation in the local speckle noise intensity. The  $J$ - and  $H$ -band spectrum of FU Ori B will be presented in a subsequent paper. This experiment explicitly demonstrates the utility of using an IFS for speckle-suppression applications.

### 3.3. Measured Contrast

PSSP provides sensitivity measurements for each star reduced. Contrast is calculated at several stages during the reduction, including with raw (pre-processed) data, high-pass Fourier-filtered data, and speckle-suppressed data. Speckle noise is found by measuring the local standard deviation of intensity variations occurring in a box of size  $5 \times 5$  PSF widths, where flux levels are calculated on the scale of a single PSF by averaging over several pixels. Speckle noise values are then divided by the stellar peak intensity measured in unocculted frames, taking into account the different integration times, to calculate the relative brightness of directly detectable companions. We also estimate the noise floor set by photon arrival statistics by



**Figure 8.** Stacked  $JH$  images of the star FU Ori before (left) and after (right) running PSSP. The companion, FU ORI B, is located within the AO control box and has a brightness comparable to the scattered-light halo of the primary. Its unambiguous detection following post-processing demonstrates the principal of using an IFS to discriminate between speckles and faint off-axis sources.

measuring the square root of the mean flux over the same regions. We nominally quote  $5\sigma$  contrast levels which correspond to a signal-to-noise ratio where companions become readily noticeable by eye.<sup>13</sup>

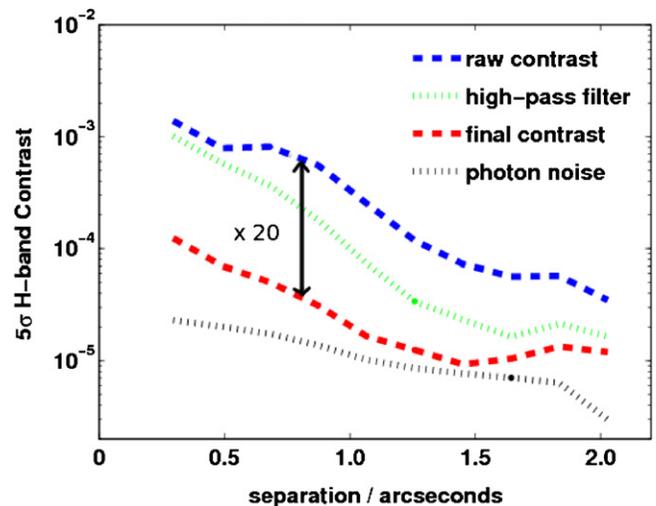
Results for the  $V = 6.7$  star HD 204277 from data taken on 2009 June 28 are shown in Figure 9. Local contrast levels using raw, Fourier-filtered, and speckle-suppressed images have been azimuthally averaged (median). We find that residual quasi-static wave-front phase errors of  $\approx 100$  nm rms limit raw (pre-processed) sensitivity to  $10^{-3}$  at the coronagraph inner-working-angle. This result is consistent with similar measurements using the PHARO instrument prior to application of the modified Gerchberg–Saxton phase retrieval algorithm (Burruss et al. 2010). High-pass filtering the data removes the low-frequency component of the stellar halo (pedestal) resulting in a factor of several improvement, most notably at separations exterior to the AO control region at  $\approx 0.6$ . Finally, chromatic differential imaging operating in tandem with the LOCI algorithm reduces remaining noise by more than an order of magnitude closer to the star at the expense of a factor of 2–3 in companion throughput. In the case of HD 204277, this results in a  $5\sigma$   $H$ -band contrast of  $2.1 \times 10^{-5}$  at  $1''$  with 20 minutes of on-source integration time.

These are the deepest contrast levels yet achieved at Palomar in the  $H$  band. Sensitivity is currently limited by calibration of non-common-path wave-front errors and instrument transmission which ultimately governs the signal-to-noise ratio of speckles. We find that the contrast continues to improve with additional exposures, but slower than a square-root relationship in time in regions very close to the star.

#### 4. SUMMARY AND CONCLUDING REMARKS

Central to the issue of generating contrast levels sufficient to directly image extrasolar planets is subtraction of (unavoidable) quasi-static speckles that arise from instrument optical aberrations. Along with the ability to characterize companions, an IFS naturally provides the capability for suppressing this dominant noise source. The recently commissioned Project 1640 instrument at Palomar is the first to use an IFS for high-contrast imaging in combination with a coronagraph and AO.

<sup>13</sup> Speckle noise follows a Rician spatial distribution, and speckles with intensity up to  $\approx 10\sigma$  may occur (Aime & Soummer 2004; Fitzgerald & Graham 2006). An IFS however can help to discriminate between unusually bright speckles and real companions by comparing the spectrum of the candidate to that of the star.



**Figure 9.** P1640 contrast levels in the  $H$  band for the star HD 204277 taken on 2009 June 28 UT. The star is occulted by the apodized pupil Lyot coronagraph. PSSP consistently provides a gain of at least one order of magnitude in sensitivity at subarcsecond separations.

(A color version of this figure is available in the online journal.)

We have written a custom program, the PSSP, to improve instrument raw sensitivity via post-processing. The program utilizes a version of chromatic differential imaging, or spectral deconvolution, to perform PSF subtraction, as was first proposed by Sparks & Ford (2002). To this, we also use the multitude of images provided by the IFS to construct PSF references via the LOCI algorithm (Lafrenière et al. 2007a).

We have quantified the degree of spatial correlation between speckle patterns in adjacent P1640 hyper-cube images as a function of color and time. We find that spectral image diversity provides PSF references that match science images at a level comparable to, if not better than, observations of a nearby calibration star, depending on overhead time and system flexure. Analysis related to image scaling factors, throughput, and signal-to-noise ratio match well the theoretical expectations and results from other speckle diversity techniques that also use LOCI, once relevant parameters are optimized. Further, we are able to perform pseudo-real-time data processing while at the observatory using the Bluedot super-computing cluster at NExScI.

We have quantified the improvement afforded by the PSSP with three different data sets. Injecting faint companions,

recovering known companions, and calculating contrast levels before and after performing speckle suppression each indicate that on-sky sensitivities consistently improve by at least one order of magnitude at angular separations within the AO control region following post-processing. At larger separations, the intensity of residual starlight approaches the limit set by photon noise. Our sensitivity currently reaches  $5\sigma$  contrast levels of  $\approx 2.1 \times 10^{-5}$  at  $1''$  in the  $H$  band with 20 minutes of integration time. This result is comparable to other experiments at Palomar and is currently limited by calibration of non-common-path errors (Burruss et al. 2010; Crepp et al. 2010). Preliminary tests using an interferometric wave-front calibration unit show promise to reduce these errors by an order of magnitude when operating in tandem with the PALM-3000 “extreme” AO system, suggesting that contrast levels of order  $\approx 10^{-7}$  inside of an arcsecond are feasible with integration times of 1–2 hr (Bouchez et al. 2010).

IFS data provide the ability to play movies by sequentially stepping through each image in the color dimension. We have performed experiments carefully comparing our ability to detect faint companions using these color movies and the PSSP. We find that movies perform remarkably well, while the PSSP generally provides an additional factor of 2–3 in effective contrast in comparison. This gain is most noticeable in close proximity to the star, just exterior to the coronagraphic spot and is largely a result of implementing the LOCI algorithm to improve PSF matching. We play color movies as part of our standard reduction package to identify companions with brightness comparable to raw speckle noise. Stepping through wavelength channels is likewise helpful with differenced images.

It is necessary to calibrate the effects of partial subtraction resulting from LOCI to accurately measure the relative intensity of a faint companion (Marois et al. 2008; Thalmann et al. 2009; Currie et al. 2010; Bowler et al. 2010; Janson et al. 2011; Marois et al. 2010a). This is nominally accomplished by injecting fake companions with known brightness into each data cube. Our photometric analysis indicates that 10% uncertainties are typical for angular separations within the AO control region and that larger uncertainties, by factors of several, may occur very close to the star where only marginally sufficient wavelength diversity is achieved.

High-contrast astrometry is likewise challenging (Digby et al. 2006). The inclusion of astrometric grid spots aids with locating the position of the occulted star, but aggressive differential imaging techniques modify the shape and flux distribution of the companion PSF, degrading precision. Our preliminary analysis using fake companions indicates that systematic errors of size  $\approx 0.5$  spatial pixels are common, even when using companion masks in reference frames. Carefully chosen reference frames can help to mitigate these effects, but 1 mas astrometry using near-Nyquist sampled data (2.1–3.4 spatial pixels per diffraction width in this case) may prove to be prohibitive unless more advanced algorithms are developed.

This paper demonstrates one of the two important benefits provided by an IFS for high-contrast imaging: the automatic generation of many useful PSF references for chromatic differential imaging and the detection of faint companions. A subsequent paper by L. Pueyo et al. (2011, in preparation) will discuss the challenges and possible solutions for accurately extracting their spectra.

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